

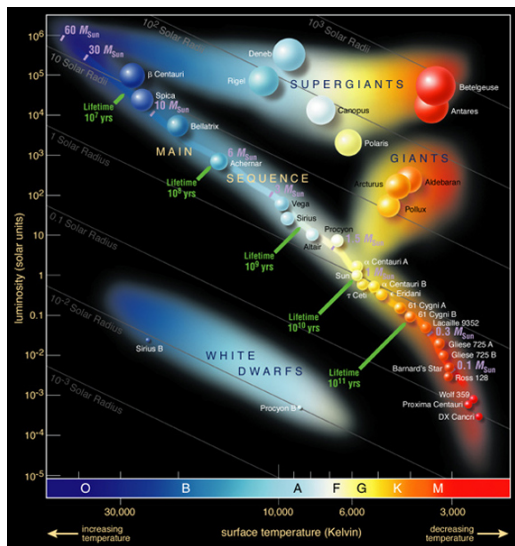
# Inversion of stellar fundamental parameters from Espadons-Narval high-resolution spectra

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- **Sources and characteristics of the data**
- **A PCA-based inversion method**
- **First results and discussion**

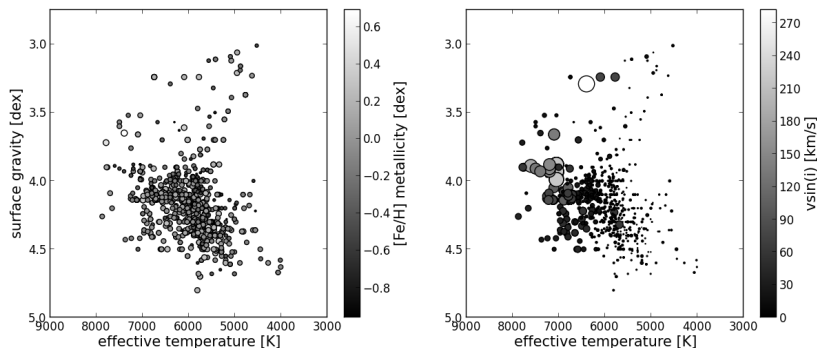
# Stellar classification



→ The Hertzsprung–Russell diagram.

- **Espadons-Narval database** ([polarbase.irap.omp.eu](http://polarbase.irap.omp.eu))
  - $\mathcal{R} \sim 65\,000$  ; spectral coverage  $\sim 380\text{-}1000$  nm
  - 180 000 independent spectra, for 2 000 distinct targets
  - Covers quite well the H–R diagram
  - see *Petit et al. (2014)* for a more complete description
- **Elodie Spectral Library** (*Prugniel et al. 2007*)
  - **Observed spectra** for the training database
  - $\mathcal{R} \sim 48\,000$  ; spectral coverage  $\sim 390\text{-}680$  nm
  - 905 spectra selected for FGK stars
- **Projected rotational velocities** ( $v \sin i$ )
  - Taken from VizieR@CDS III/244 (*Głebocki & Gnaciński 2005*)
- **Harvesting data:** see e.g., *Paletou & Zolotukhin (2014)*  
[arXiv:1408.7026]

# The training database



→ Graphical summary of the coverage in stellar parameters corresponding to the content of our Elodie spectra training database (for FGK stars).

- **Variance-covariance matrix** ( $\{S_i\}$ : training database)

$$\mathbf{C}_{(N_\lambda; N_\lambda)} = (\mathbf{S} - \bar{\mathbf{S}})^T \cdot (\mathbf{S} - \bar{\mathbf{S}}) ; N_\lambda = 8000 \quad (1)$$

- Knowing its **eigenvectors**  $\vec{e}_k$  ( $k \leq k_{max}$ ) allows for the computation of **projection coefficients**:

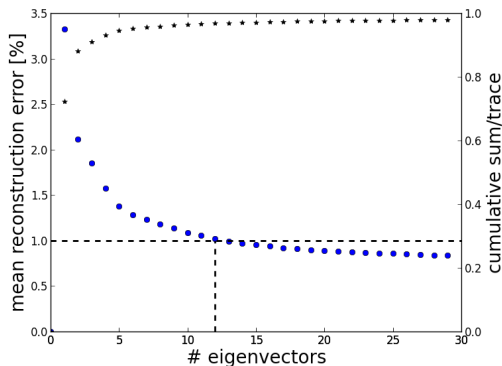
$$p_{jk} = (\vec{S}_j - \bar{\mathbf{S}}) \cdot \vec{e}_k \quad (2)$$

- **Reduction of dimensionality**

– Each spectra of  $N_\lambda$  (8000) points will be described by “just”  $k_{max}$  ( $\ll N_\lambda$ ) projection coefficients

- **How many orders  $k_{max}$  to keep?**

# Reduction of dimensionality



→ Reconstruction error  $E(k_{\max}) = \left\langle \left( \frac{|\bar{S} + \sum_{k=1}^{k_{\max}} p_{jk} \vec{e}_k - S_j|}{S_j} \right) \right\rangle$  as a function of the number of eigenvectors used for the computation of the projection coefficients  $p$ .

# Inversion of stellar parameters

- Nearest neighbour(s) search
- Observed spectra:  $\vec{O}(\lambda)$
- Projected on  $k_{max}$  eigenvectors:  $\varrho_k = (\vec{O} - \vec{S}) \cdot \vec{e}_k$
- Compute squared (euclidian) distances:

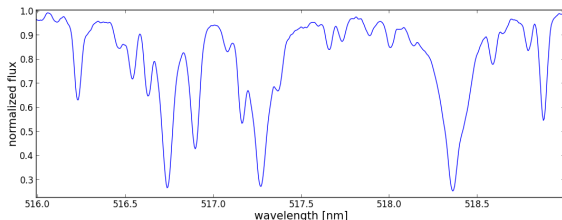
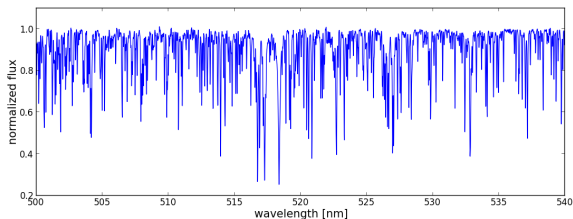
$$d_j^{(O)} = \sum_{k=1}^{k_{max}} (\varrho_k - p_{jk})^2 \quad (3)$$

- Consider neighbours in domain:

$$\min(d_j^{(O)}) \leq d_j^{(O)} \leq 1.2 \times \min(d_j^{(O)}) \quad (4)$$

- Simple mean of every stellar parameters of the identified neighbours  $\rightarrow (T_{\text{eff}} ; \log g ; [\text{Fe}/\text{H}] ; v \sin i)$

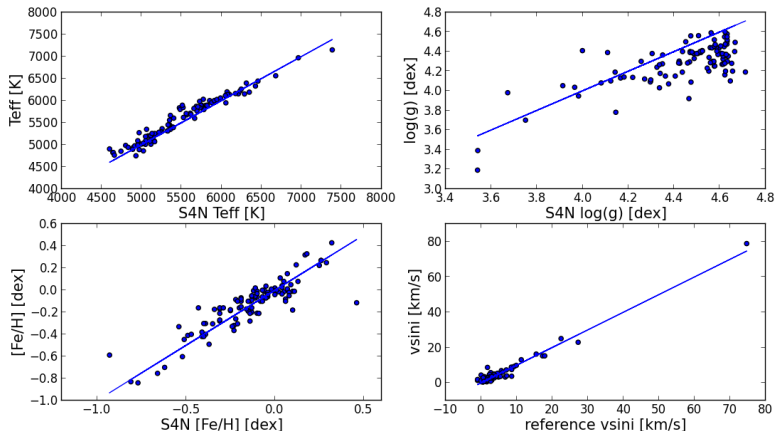
# Which spectral domain(s)?



→ We chose to test our approach at the vicinity of the Mg I b triplet (known to be  $\log g$  sensitive).



# Inverting $S^4N$ data



→ First tests by inverting data from the  $S^4N$  survey (*Allende Prieto et al. 2004*).

# Inverting PolarBase data

- A selection of **140 spectra (of FGK stars)** in common with the “Spectroscopic Properties of Cool Stars” data collection (*Valenti & Fischer 2005*)

Parameter	$T_{\text{eff}}$	$\log g$	[Fe/H]	$v \sin i$
bias	23	0.10	0.02	-0.04
$\sigma$	115	0.19	0.10	1.68

- **Advantages of working with observed spectra**
  - **Object to objects relation!**
  - **Independent** from whatever method of determination of stellar parameters have been used for the reference spectra
  - One may **also** invert **any other parameter** contributing to the spectral signature of the star

- **PCA is simple to implement**
- **It leads to a fast and robust inversion method**
- **Very encouraging results with Espadons and Narval spectra:**  
further details in *Paletou et al. 2014 (A&A, in press)*
- **Future work is...**
  - Use (also) synthetic spectra, for spectral domains not covered by Elodie like e.g., the **RVS@GAIA Ca II Infra-Red Triplet domain**
  - Combine inversion results from several spectral domains
  - Extend the range of application to **A–FGK–M** spectral types
  - Implement at [polarbase.irap.omp.eu](http://polarbase.irap.omp.eu)
  - ...